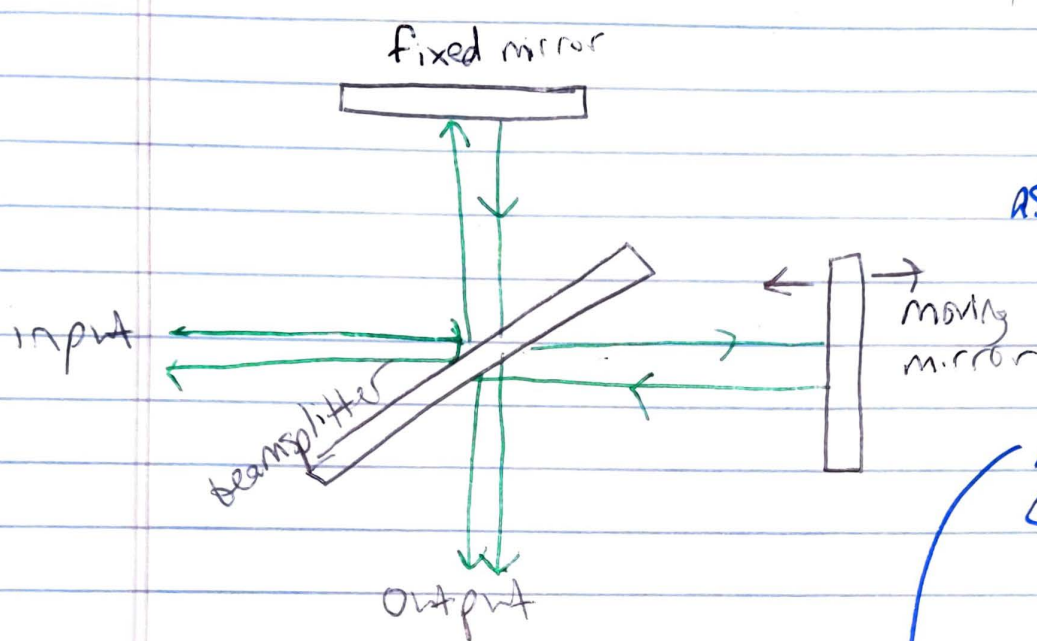


# Fourier Transform Spectroscopy



output

$$E = h_1 e^{i\phi_1} + h_2 e^{i\phi_2}$$

assume  $h_1 = h_2 = h$

$$I = EE^*$$

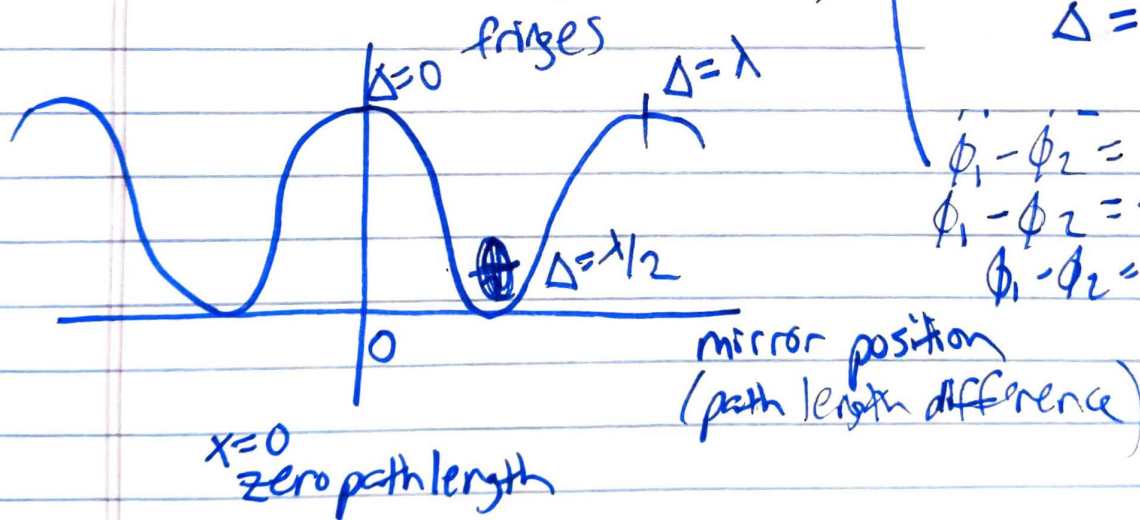
$$I = 2h^2(1 + \cos(\phi_1 - \phi_2))$$

$$I = 2h^2(1 + \cos(\frac{2\pi\Delta}{\lambda}))$$

$$\Delta = \lambda(\phi_1 - \phi_2) / 2\pi$$

$\Delta =$  path length difference

Monochromatic light (laser)



move mirror at  $v$

$$\Delta = vt$$

$$\phi_1 - \phi_2 = 2\pi \Delta / \lambda$$

$$\phi_1 - \phi_2 = 2\pi vt / \lambda$$

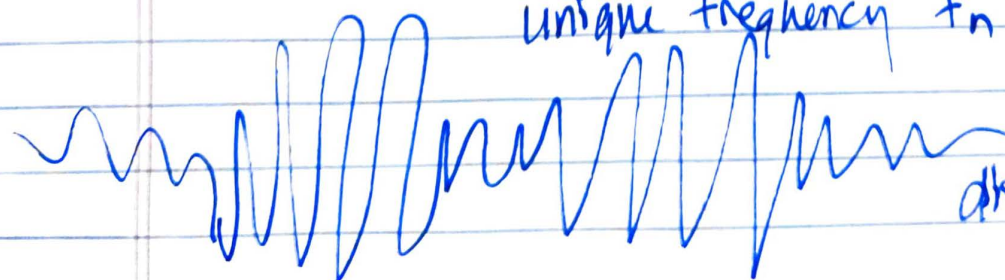
$$\phi_1 - \phi_2 = 2\pi ft$$

$$f = \frac{2v}{\lambda}$$

Broadband Light

$$I(t) = \sum_n 2h_n^2 \cos(2\pi f_n t)$$

unique frequency  $f_n$  for each wavelength



can invert w/  
discrete Fourier Transform

"Fellgett advantage"  
"multiplex advantage"

In limit of detector noise dominated measurement,  
noise is not increased by measuring many wavelengths  
S/N is improved by  $\sqrt{\#}$  of spectral elements

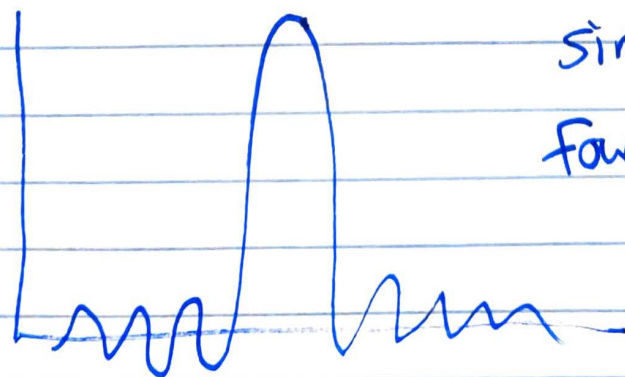
If shot noise / photon noise dominates,  
advantage disappears  
noise  $\propto \sqrt{P} \propto \sqrt{\#}$  of spectral elements

Hard limit on performance set by path-length difference  
and step size

step-size sets highest frequency / shortest wavelength

spectral resolution is set by largest possible excursion /  
path length difference

What does an unresolved spectral line look like?

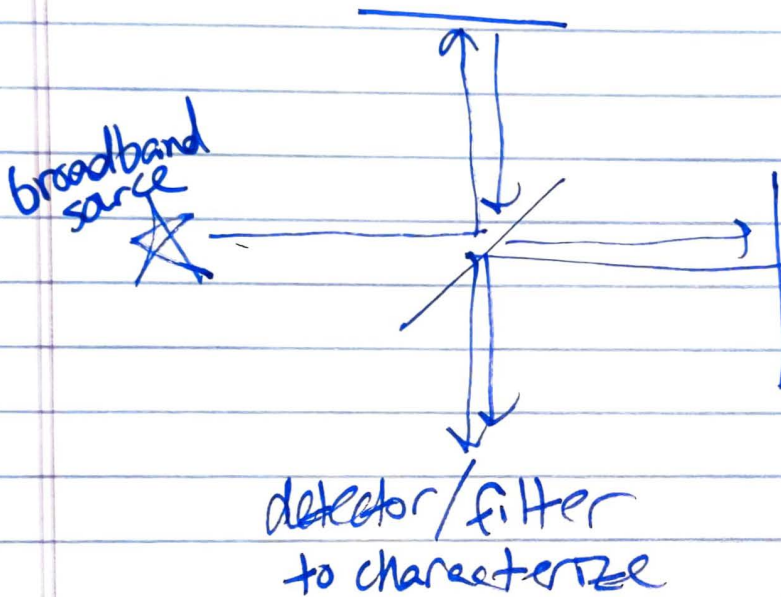


sinc function w/ ringing

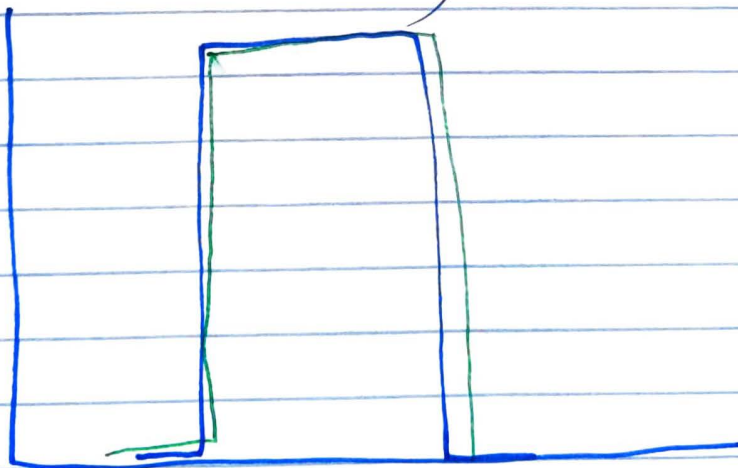
Fourier series is truncated

fix by apertization

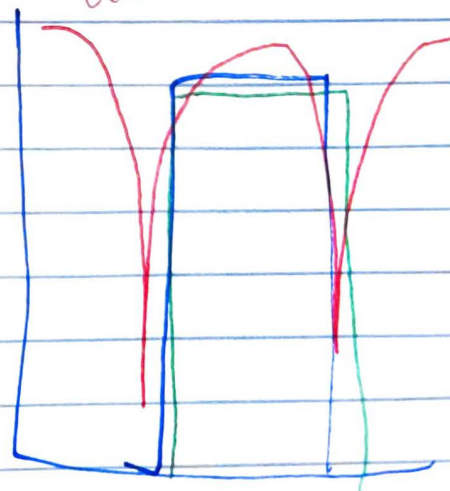
FTS can be used for observation  
or  
calibration



Bandpass matching

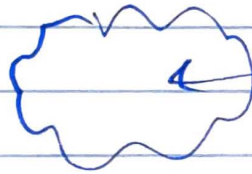
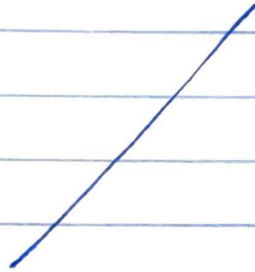


atmospheric absorption

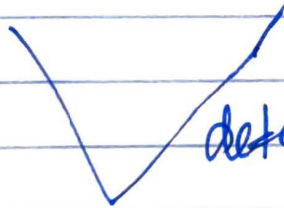


FTS can be used to measure absorption

Sun / Blackbody



← sample



detector